

LETTER TO THE EDITOR

# An ALMA view of the interstellar medium of the $z = 4.77$ lensed starburst SPT-S J213242-5802.9

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## ABSTRACT

We present ALMA detections of the [NII] 205  $\mu\text{m}$  and CO(12–11) emission lines, and the tentative detection of [CI]  $^3\text{P}_1$ – $^3\text{P}_0$  for the strongly lensed ( $\mu = 5.7 \pm 0.5$ ) dusty, star-forming galaxy SPT-S J213242-5802.9 (hereafter SPT2132-58) at  $z = 4.77$ . The [NII] and CO(12–11) lines are detected at  $11.5$  and  $8.5\sigma$  levels, respectively, by our band 6 observations. The [CI] line is detected at  $3.2\sigma$  after a reanalysis of existing band 3 data. The [CI] luminosity implies a gas mass of  $(3.8 \pm 1.2) \times 10^{10} M_\odot$ , and, consequently, a very short depletion timescale of  $34 \pm 13$  Myr and a CO luminosity to gas mass conversion factor  $\alpha_{\text{CO}}$  of  $1.0 \pm 0.3 M_\odot (\text{K km s}^{-1} \text{pc}^2)^{-1}$ . SPT2132-58 is an extreme starburst with an intrinsic star formation rate of  $1100 \pm 200 M_\odot/\text{yr}$ . We find a [CII]/[NII] ratio of  $26 \pm 6$ , which is the highest ratio reported at  $z > 4$ . This suggests that SPT2132-58 hosts an evolved interstellar medium ( $0.5 Z_\odot < Z < 1.5 Z_\odot$ ), which may be dominated by photodissociation regions. The CO(2–1) and CO(5–4) transitions have lower CO to far-infrared ratios than local and high-redshift samples, while CO(12–11) is similar to these samples, suggesting the presence of an additional very excited component or an active galactic nucleus.

**Key words.** galaxies: starburst – galaxies: ISM – galaxies: high-redshift – galaxies: star formation – submillimeter: galaxies

## 1. Introduction

The presence of a large population of massive, dusty galaxies hosting high star formation rates ( $SFR > 500 M_\odot/\text{yr}$ ) at high redshift (e.g. Smail et al. 1997; Blain et al. 2002; Chapman et al. 2005) cannot be explained by the standard semi-analytical models of galaxy formation and hydrodynamical simulations (e.g. Cousin et al. 2015; Sparre et al. 2015). Recently, several of these extreme star-forming galaxies have been found at  $z > 4$  (e.g. Riechers et al. 2013; Vieira et al. 2013; Dowell et al. 2014; Smolčić et al. 2015), indicating extremely quick assembly of these structures. Their high SFRs suggest the presence of large gas reservoirs and their large dust content indicates that their interstellar medium (ISM) is metal enriched. This raises the question of how such a massive and mature ISM can be built so early?

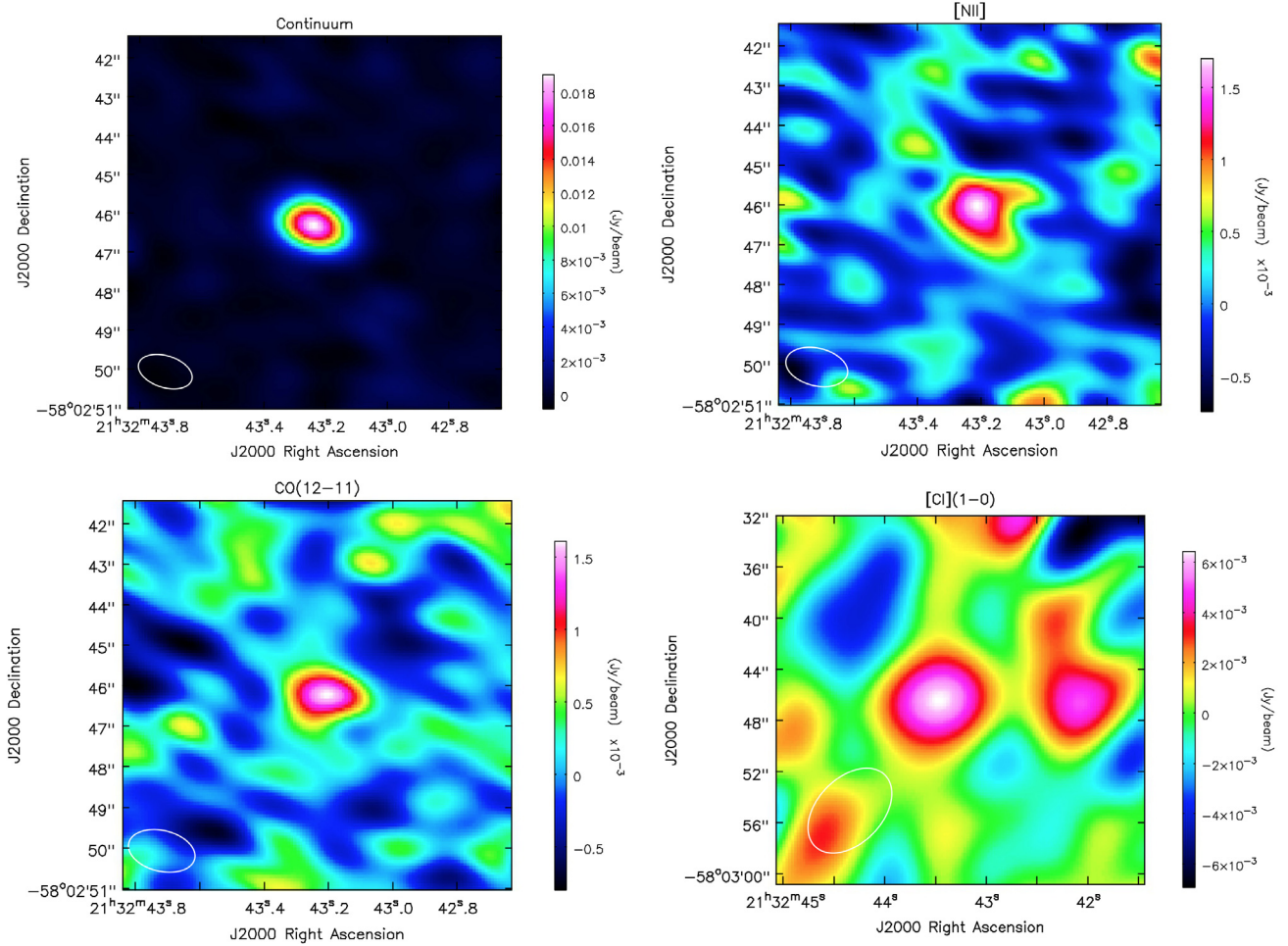
Studying the ISM of these  $z > 4$  galaxies is difficult using the rest-frame optical because of the strong dust attenuation (da Cunha et al. 2015). The dust continuum and the (sub-)millimeter lines from the ISM are thus the main information we can obtain and the continuum and [CII] can now be

routinely detected by ALMA (e.g., Scoville et al. 2014; Capak et al. 2015). The [CII]/[NII] line ratio was proposed as a new diagnostic of the ISM in star-forming galaxies (Nagao et al. 2012), but its use is hampered by the faintness of the [NII] line requiring long integration times, even with ALMA (Nagao et al. 2012; Decarli et al. 2014). This problem can be solved by targeting sources amplified by gravitational lensing. The sample of strongly lensed sources found by the South Pole Telescope (SPT; Carlstrom et al. 2011) is particularly suited for these follow-up studies because of its high median redshift ( $\langle z \rangle = 3.9$ ; Strandet et al. 2016).

In this Letter, we report the detections of the [NII], CO(12–11), and [CI]  $^3\text{P}_1$ – $^3\text{P}_0$  lines in SPT-S J213242-5802.9 at  $z = 4.768$  (hereafter, SPT2132-58; Weiß et al. 2013; Gullberg et al. 2015). We assume the Planck Collaboration XIII (2015) cosmology ( $H_0 = 67.8 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_m = 0.308$ ).

## 2. Data

SPT2132-58 was observed with ALMA in band 6 on February 23 2014 (2012.1.00994.S, PI: De Breuck) with a



**Fig. 1.** Maps of the continuum (*top left*), the [NII] line (*top right*), the CO(12–11) line (*bottom left*), and the [CI](1–0) line (*bottom right*). The size of the beam is indicated in the lower left corner. [NII] and CO(12–11) were integrated between  $-300$  and  $+300$   $\text{km s}^{-1}$ . Since the [CI](1–0) line is narrower and fainter, it was integrated between  $-100$  and  $+100$   $\text{km s}^{-1}$  to maximize the signal-to-noise ratio. The beam in band 3 ([CI](1–0)) is much larger than in band 6. We thus stretch the  $x$ - and  $y$ -axis by a factor 3.

relatively compact configuration (minimal baseline = 15 m, max = 399 m, rms = 140 m) and good weather (PWV = 0.9 mm). The on-source integration time is 470 s. The data were calibrated by the standard ALMA pipeline. The four spectral windows are centred on 237.6, 239.5, 252.9, and 254.8 GHz. Each spectral window has 480 channels and covers 1875 MHz. We produced the images of our source with CASA (McMullin et al. 2007) using the `clean` routine, and we chose a natural weighting to optimize the sensitivity (see Sect. 3). This results in a synthesized beam size of  $1.55'' \times 0.91''$ . The typical RMS noise is 0.36 mJy/beam for the continuum (combining the four spectral windows) and 1.6 mJy/beam/(25  $\text{km s}^{-1}$  channel) for the spectra. The spectra were extracted using an elongated aperture twice as large in diameter than the beam to avoid losing flux because our source is marginally resolved. We also produced a continuum image with a Briggs weighting with `robust` = 0.5, since the signal-to-noise ratio is already very high ( $\text{SNR} = 52$ ; see Fig. 1).

In addition, we reanalyzed band 3 data obtained during the cycle 0 (Weiß et al. 2013) using the same approach (see Fig. 1). We used a natural weighting instead of a Briggs weighting with `robust` = 0.5. The beam size is  $7.7'' \times 4.9''$  and the noise is 3.2 mJy/beam/(50  $\text{km s}^{-1}$  channel). The natural weighting reduces the noise in the spectrum and improves the signal-to-noise ratio of the [CI] line from 2.9 to 3.2.

### 3. Results

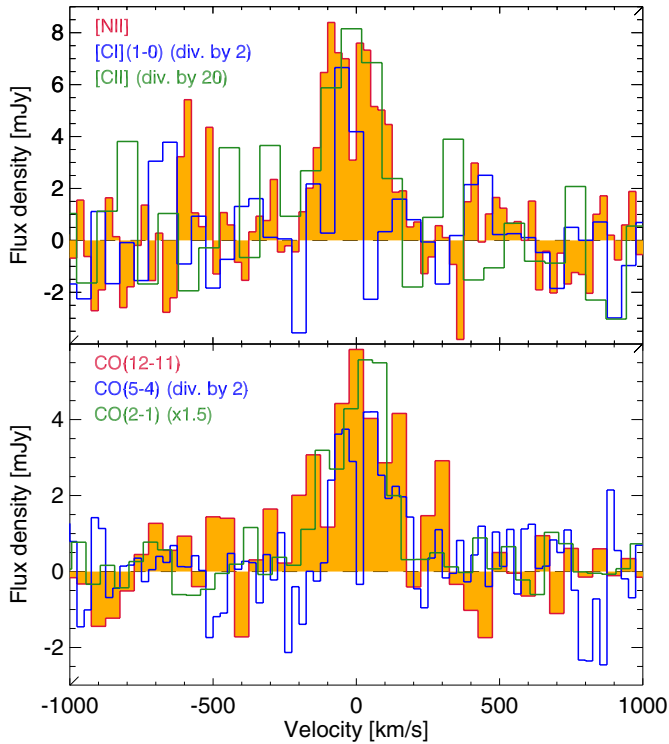
The [NII] and CO(12–11) lines are detected unambiguously at  $11.5$  and  $8.5\sigma$ , respectively (see Fig. 2). The reanalysis of the cycle 0 data allowed us to obtain a tentative detection of [CI] at  $3.2\sigma$ . The fluxes of the lines are listed in Table 1. Finally, we measured the flux density of the source and found an integrated flux  $26.4 \pm 0.5$  mJy at 246 GHz. The [NII] and CO(12–11) lines were not excluded from this measurement, since ancillary bolometric data also include lines and their contribution is only 1.3% of the total continuum flux integrated over the four spectral windows. The source is marginally resolved along its major axis. The AIPS and CASA softwares provide a similar best estimate of the source size of  $(0.6 \pm 0.2)''$ . This value is consistent with the size in the image plane predicted by our best lens model ( $0.67''$ ; Spilker et al., in prep.).

We compared the profiles of these new detections with the previously reported detections of CO(5–4) by Weiß et al. (2013), [CII] by Gullberg et al. (2015), and CO(2–1) by Aravena et al. (2016). All the lines are compatible with a single Gaussian function (highest reduced  $\chi^2$  is 1.4, which is compatible at  $1\sigma$  with unity considering their number of degrees of freedom). The results of the fits are given in Table 1. The [NII], [CII], CO(5–4), and CO(2–1) lines have similar full widths at half maximum

**Table 1.** Summary of the properties of the detected lines of SPT2132-58.

Line	$\nu_{\text{rest}}$ GHz	$I_{\text{line}}^a$ Jy km s $^{-1}$	$\mu L'_{\text{line}}$ $10^{10}$ K km $^{-1}$ pc $^2$	$\mu L_{\text{line}}$ $10^8 L_{\odot}$	$FWHM$ km s $^{-1}$	Instrument	Reference
CO(2-1)	230.54	$0.85 \pm 0.07$	$18.38 \pm 1.51$	$0.71 \pm 0.06$	$225 \pm 17$	ATCA	Aravena et al. (2016)
[CI](1-0)	492.16	$0.86 \pm 0.27$	$4.07 \pm 1.29$	$1.54 \pm 0.49$	...	ALMA band 3	This work
CO(5-4)	576.35	$1.41 \pm 0.26$	$4.88 \pm 0.90$	$2.96 \pm 0.54$	$231 \pm 70$	ALMA band 3	Weiß et al. (2013)
CO(12-11)	1383.24	$1.36 \pm 0.16$	$0.82 \pm 0.10$	$6.84 \pm 0.81$	$331 \pm 78$	ALMA band 6	This work
[NII]	1461.10	$1.73 \pm 0.15$	$0.93 \pm 0.08$	$9.20 \pm 0.79$	$240 \pm 41$	ALMA band 6	This work
[CII]	1900.54	$34.90 \pm 6.90$	$11.10 \pm 2.19$	$241.53 \pm 47.75$	$212 \pm 43$	APEX/FLASH	Gullberg et al. (2015)

**Notes.** These apparent luminosities have not been corrected for lensing magnification. <sup>(a)</sup> The integrated line flux  $I_{\text{CO}}$  is computed between  $-1.5$  and  $1.5 FWHM$ .



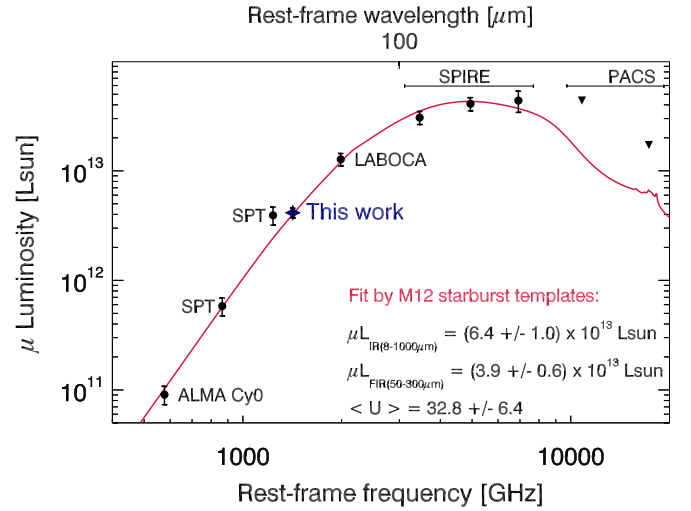
**Fig. 2.** Upper panel: profile of the fine structure lines of SPT2132-58 ([NII] in orange, [CI](1-0) in blue, and [CII] in green). Lower panel: profile of the CO lines (12-11 in orange, 5-4 in blue, and 2-1 in green).

(FWHMs). The CO(12-11) line is slightly wider than the other lines ( $1.0\sigma$  compared to [NII]). The low SNR of [CI] prohibits us from determining the FWHM.

## 4. Discussion

### 4.1. Star formation efficiency and nature of the source

SPT2132-58 is gravitationally magnified by another galaxy on the line of sight. Our best model, based on the ALMA cycle-0, high-resolution follow-up of the source, estimates that the magnification is  $\mu = 5.7 \pm 0.5$  (Spilker et al., in prep.). We determined its apparent total infrared luminosity ( $\mu L_{\text{IR}(8-1000\mu\text{m})} = (6.4 \pm 1.0) \times 10^{13} L_{\odot}$ ) and the far-infrared luminosity ( $\mu L_{\text{FIR}(50-300\mu\text{m})} = (3.9 \pm 0.6) \times 10^{13} L_{\odot}$ ) by fitting the continuum measurements made by ALMA band 3, SPT, LABOCA, *Herschel*/SPIRE (Strandet et al., in prep.), and ALMA band 6 (Sect. 3 and Fig. 3). We used the starburst templates of Magdis et al. (2012) to fit the continuum. The quality of the fit is good ( $\chi^2_{\text{red}} = 1.1$ ), but the SPT point at 220 GHz exhibits a  $2\sigma$  excess. This tension

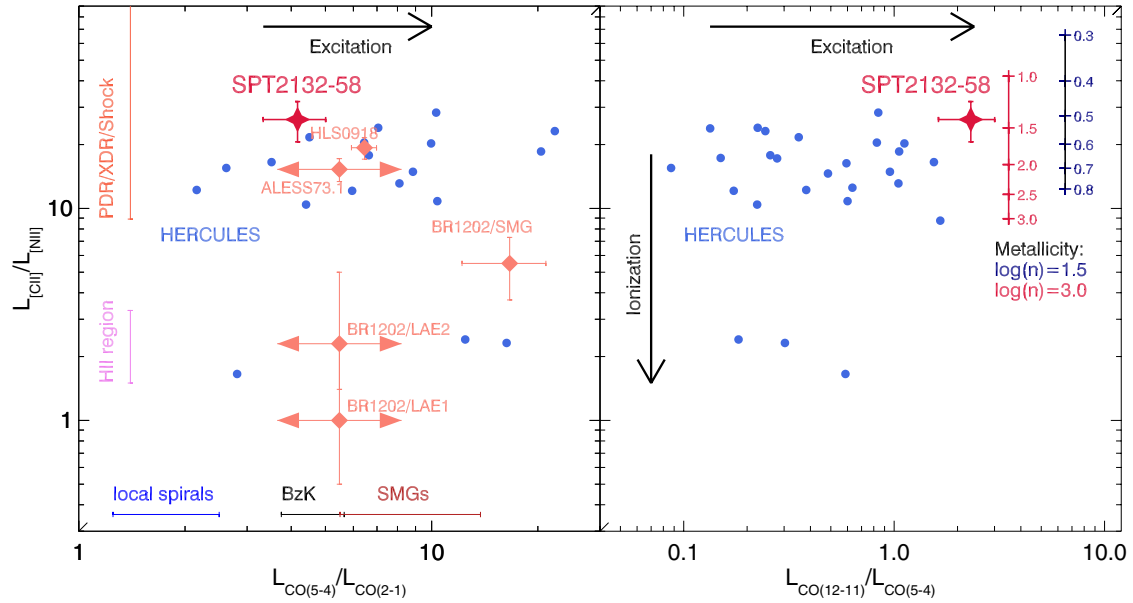


**Fig. 3.** SED of SPT2132-58, including a compilation of measurements from Strandet et al. (in prep.), (filled black circles for detections, downwards triangles for  $3\sigma$  upper limits) and our new ALMA Cycle 1 continuum measurement (blue filled star), and best fit by the starburst templates of the Magdis et al. (2012) model (red line). The  $\langle U \rangle$  parameter is the mean intensity of the radiation field heating the dust compared to the solar neighbourhood.

could be explained by a peak of noise at the position of our source and not by an astrophysical contaminant since no excess is found in the other SPT band. We found an apparent luminosity  $\mu L_{\text{IR}}$  of  $(6.4 \pm 1.0) \times 10^{13} L_{\odot}$ . Because of the high dust temperature of the source ( $T_{\text{dust}} = 41$  K, Aravena et al. 2016), the CMB has a negligible impact on this measurement (da Cunha et al. 2013). The intrinsic  $L_{\text{IR}}$  is thus around  $10^{13} L_{\odot}$  putting this source at the edge of the hyper-luminous infrared galaxy (HyLIRG;  $L_{\text{IR}} > 10^{13} L_{\odot}$ ) category. This corresponds to an intrinsic SFR of  $1120 \pm 200 M_{\odot}/\text{yr}$  (Kennicutt 1998 assuming a Chabrier 2003 initial mass function). This SFR range is expected to be populated only by episodic starbursts, according to the Béthermin et al. (2012) model.

We estimated the gas mass from the [CI] emission (Papadopoulos & Greve 2004). Assuming the abundance and excitation factors of Alaghband-Zadeh et al. (2013), we found an apparent gas mass  $\mu M_{\text{gas}}$  of  $(21.5 \pm 6.7) \times 10^{10} M_{\odot}$  and a demagnified gas mass of  $(3.8 \pm 1.2) \times 10^{10} M_{\odot}$ . A systematic analysis of all [CI] lines detected in SPT sources and a detailed discussion about the accuracy of [CI]-derived gas mass will be presented in Bothwell et al. (in prep.). We can also estimate the  $\alpha_{\text{CO}}$  of this object from the  $L'_{\text{CO}(2-1)}$  line luminosity measured by Aravena et al. (2016), converting it into  $L'_{\text{CO}(1-0)}$  using the mean SMG





**Fig. 4.** CO excitation (CO(5–4)/CO(2–1) in the *left panel* and CO(12–11)/CO(5–4) in the *right panel*) versus [CII]/[NII] luminosity ratio (the ratio reported for ALESS73.1 is computed using the updated [CII] flux of De Breuck et al. 2014). Red star: SPT2132-58. Salmon diamonds: other high- $z$  sources from the literature (see text). Blue circles: HERCULES local sample (Rosenberg et al. 2015). [NII] luminosities are from Kamenetzky et al. (2015). The left and right arrows indicate that this source does not have available measurement of its CO ratio. The range of CO(5–4)/CO(2–1) ratio for various samples is indicated in the bottom of the left panel (Daddi et al. 2015). The range of ratio expected for HII and PDR/XDR/shock dominated regions from Decarli et al. (2014) are indicated on the left. The metallicity estimated from [CII]/[NII], following the Nagao et al. (2012) method, is indicated on the right for  $\log(n) = 1.5$  (navy) and  $\log(n) = 3.0$  (blue). We assume  $\log(U_{\text{HII}}) = -3.5$  for both.

ratio (Carilli & Walter 2013), and assuming the gas mass provided by [CI]. We find  $\alpha_{\text{CO}} = 0.99 \pm 0.32 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ . This agrees with  $\alpha_{\text{CO}}$  of  $0.9 \pm 0.3 (\text{K km s}^{-1} \text{pc}^2)^{-1}$  derived by Aravena et al. (2016) from the dust mass assuming a gas-to-dust mass ratio of 100. This value is compatible with the classical  $\sim 0.8 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$  found for local compact starbursts (Downes & Solomon 1998).

The depletion timescale associated with this gas mass and SFR is  $34 \pm 13 \text{ Myr}$ . This is a factor of 3 shorter than typical strong starbursts and more than an order of magnitude less than normal star-forming galaxies at this redshift (e.g. Béthermin et al. 2015). The short depletion timescale indicates that this source is an extreme starburst with a moderate gas content, but a very high star formation efficiency. The depletion timescale could however be slightly underestimated if this source is observed at the end of the starburst. In this case, the SFR would be quickly decreasing. The SFR derived from  $L_{\text{IR}}$  would thus be lower than the actual SFR, since this tracer is sensitive to the SFR in the last 10–100 Myr (Kennicutt 1998).

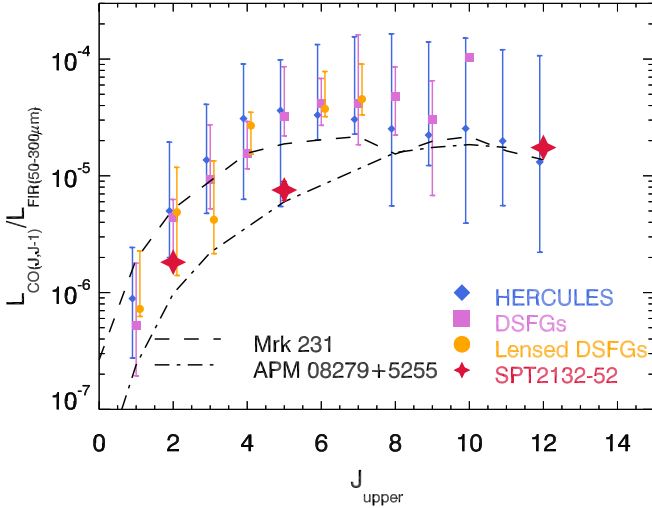
#### 4.2. Properties of ISM

SPT2132-58 is one of the few high-redshift sources that have been detected in both [CII] and [NII]  $205 \mu\text{m}$  ( $y$ -axis of Fig. 4). The previous detections were obtained in HLS J091828.6+514223 (Combes et al. 2012; Rawle et al. 2014), ALESS73.1 (Nagao et al. 2012), and one SMG and two LAEs (tentative detections) in the BR1202-0725 system (Decarli et al. 2014). We find a [CII]/[NII] luminosity ratio of  $26.3 \pm 5.7$  in SPT2132-58, which is the highest reported so far at high redshift. HLS J0918 has a similar ratio within  $1\sigma$ , but ALESS73.1 is two times lower. The members of the BR1202-0725 system have a much lower ratio (4–26 times). Finally, we compared SPT2132-58 with the HERCULES sample of local (U)LIRGs (Rosenberg et al. 2015). All but three of the 29 HERCULES sources have a

lower ratio. However, the [CII]/[NII] ratio of SPT2132-58 is not uncommon in the local Universe and is close to, e.g. the mean value of 30 found by Spinoglio et al. (2015) in local Seyfert galaxies.

The [CII]/[NII] ratio of SPT2132-58 corresponds to a PDR/XDR/shock-dominated regime according to Decarli et al. (2014). For galaxies in this regime, the [CII] originates from both HII regions and external layers of PDRs, while [NII] only comes from HII regions. Using a model that takes both PDR and HII regions into account, Nagao et al. (2012) showed that the observed [CII]/[NII] ratio depends on metallicity. This occurs since the PDR/HII fraction of [CII] varies with metallicity. Other parameters affecting this ratio include density and, to a lesser extent, the radiation field. Figure 4 shows the predictions of Nagao et al. (2012) for both a low (navy) and a high (red) density, but only for a single ionization parameter  $\log(U_{\text{HII}}) = -3.5$ . The [CII]/[NII] ratio of SPT2132-58 corresponds to a metallicity in the 0.5 to  $1.5 Z_{\odot}$  range. While we could have expected a very low metallicity at such a high redshift, our source is not far from solar metallicity. Nevertheless, our source has a slightly higher [CII]/[NII] ratio than the HERCULES sample. Although the [CII]/[NII] ratio depends on several other physical parameters (e.g. density), the higher ratio found in SPT2132-58 suggests a slightly lower metallicity than in local starbursts.

We also compared the CO excitation with the samples cited previously ( $x$ -axis of Fig. 4). SPT2132-58 has a CO(5–4)/CO(2–1) ratio close to the average of the HERCULES sample. Assuming  $T_{\text{kin}} = 40 \text{ K}$ , the CMB at  $z = 4.77$  has only an impact of 15% on this ratio (da Cunha et al. 2013). The CO(5–4)/CO(2–1) ratio lies at the border between the value found for BzK galaxies and SMGs (Daddi et al. 2015) and is thus standard for a high-redshift galaxy. However, these two lines have a lower  $L_{\text{CO}}/L_{\text{FIR}}(50\text{--}300 \mu\text{m})$  ratio than the other samples of DSFGs by a factor of 2 and 3 for the CO(2–1) and CO(5–4) transitions, respectively (see Fig. 5). This is compatible with the



**Fig. 5.** CO spectral line energy distribution (SLED) of SPT2132-58 normalized by  $L_{\text{FIR}}$ . Our results are compared with three data compilations of Greve et al. (2014; the error bars indicate the lowest and highest value of their samples). The dashed line is the SLED of Mrk231 (Rosenberg et al. 2015) and the dash-dotted line is the SLED of APM 08279 +5255 (Downes et al. 1999; Weiß et al. 2007; Riechers et al. 2009; Bradford et al. 2011).

moderate gas content and high SFR found in Sect. 4.1, since  $L_{\text{CO}}$  correlates with the gas content and SFR with  $L_{\text{FIR}(50-300\mu\text{m})}$ . The CO(5–4) line is also a factor of  $3.2 \pm 0.4$  lower than the average CO(5–4) flux of the SPT sample measured by stacking in Spilker et al. (2014).

In contrast, the  $L'_{\text{CO}(12-11)}/L_{\text{FIR}(50-300\mu\text{m})}$  ratio is similar to the HERCULES sample, but the CO(12–11)/CO(5–4) ratio of SPT2132-58 is higher than all of the objects of the HERCULES sample. There is thus a deficit of flux of the low- and mid- $J$  CO lines, but a strong CO(12–11) line that may be emitted by different regions and/or excited by another mechanism. This could involve an active galactic nucleus (AGN, e.g. Gallerani et al. 2014) and/or regions with a particularly high radiation field. This effect could possibly be amplified by differential magnification (Serjeant 2012). If a very excited compact region is located on a caustic and has its flux magnified by a much larger factor than the rest of the system, the apparent ratio between a high- $J$  and a low- $J$  CO line can be higher than the intrinsic ratio. In Fig. 5, we compare the SLED of SPT2132-58 with the AGN-host galaxies Mrk231 and APM 08279+5255. The SPT2132-58 has a similar  $L_{\text{CO}(12-11)}/L_{\text{FIR}(50-300\mu\text{m})}$  ratio to these two objects. At  $J \leq 5$ , there is a factor of 5 difference between Mrk231, which matches the properties of the HERCULES sample, and APM 08279+5255, which is a strong outlier. SPT2132-58 has intermediate properties and its SLED does not allow us to reach a conclusion about the presence or not of an AGN. Future [OI] or mid-infrared observations are required to test the hypothesis.

## 5. Conclusion

We investigated the physical properties of SPT2132-58, a lensed source at  $z = 4.77$ . We determined that this source is an HyLIRG ( $\text{SFR} = 1120 \pm 200 M_{\odot}/\text{yr}$ ) with a relatively modest gas reservoir ( $3.8 \pm 1.2 \times 10^{10} M_{\odot}$ ) and a very short gas depletion timescale ( $34 \pm 13$  Myr), indicating a very intense starburst. The  $\alpha_{\text{CO}}$  conversion factor ( $0.99 \pm 0.32$ ) is similar with that measured in starbursts at lower redshift. The [CII]/[NII] ratio ( $26.3 \pm 5.7$ )

is the highest found in a high-redshift galaxy and indicates a chemically-evolved ISM ( $0.5 Z_{\odot} < Z < 1.5 Z_{\odot}$ ). This source is an important laboratory to understand the extremely quick stellar mass assembly of the first massive galaxies.

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## References

- Alaghband-Zadeh, S., Chapman, S. C., Swinbank, A. M., et al. 2013, *MNRAS*, **435**, 1493
- Aravena, M., Spilker, J. S., & et al. 2016, *ApJ*, submitted
- Béthermin, M., Daddi, E., Magdis, G., et al. 2012, *ApJ*, **757**, L23
- Béthermin, M., Daddi, E., Magdis, G., et al. 2015, *A&A*, **573**, A113
- Blain, A. W., Smail, I., Ivison, R. J., Kneib, J.-P., & Frayer, D. T. 2002, *Phys. Rep.*, **369**, 111
- Bradford, C. M., Bolatto, A. D., Maloney, P. R., et al. 2011, *ApJ*, **741**, L37
- Capak, P. L., Carilli, C., Jones, G., et al. 2015, *Nature*, **522**, 455
- Carilli, C. L., & Walter, F. 2013, *ARA&A*, **51**, 105
- Carlstrom, J. E., Ade, P. A. R., Aird, K. A., et al. 2011, *PASP*, **123**, 568
- Chabrier, G. 2003, *PASP*, **115**, 763
- Chapman, S. C., Blain, A. W., Smail, I., & Ivison, R. J. 2005, *ApJ*, **622**, 772
- Combes, F., Rex, M., Rawle, T. D., et al. 2012, *A&A*, **538**, L4
- Cousin, M., Lagache, G., Béthermin, M., Blaizot, J., & Guiderdoni, B. 2015, *A&A*, **575**, A32
- da Cunha, E., Groves, B., Walter, F., et al. 2013, *ApJ*, **766**, 13
- da Cunha, E., Walter, F., Smail, I. R., et al. 2015, *ApJ*, **806**, 110
- Daddi, E., Dannerbauer, H., Liu, D., et al. 2015, *A&A*, **577**, A46
- De Breuck, C., Williams, R. J., Swinbank, M., et al. 2014, *A&A*, **565**, A59
- Decarli, R., Walter, F., Carilli, C., et al. 2014, *ApJ*, **782**, L17
- Dowell, C. D., Conley, A., Glenn, J., et al. 2014, *ApJ*, **780**, 75
- Downes, D., & Solomon, P. M. 1998, *ApJ*, **507**, 615
- Downes, D., Neri, R., Wiklind, T., Wilner, D. J., & Shaver, P. A. 1999, *ApJ*, **513**, L1
- Gallerani, S., Ferrara, A., Neri, R., & Maiolino, R. 2014, *MNRAS*, **445**, 2848
- Greve, T. R., Leonidaki, I., Xilouris, E. M., et al. 2014, *ApJ*, **794**, 142
- Gullberg, B., De Breuck, C., Vieira, J. D., et al. 2015, *MNRAS*, **449**, 2883
- Kamenetzky, J., Rangwala, N., Glenn, J., Maloney, P. R., & Conley, A. 2015, *ApJ*, submitted [[arXiv:1508.05102](https://arxiv.org/abs/1508.05102)]
- Kennicutt, Jr., R. C. 1998, *ApJ*, **498**, 541
- Magdis, G. E., Daddi, E., Béthermin, M., et al. 2012, *ApJ*, **760**, 6
- McMullin, J. P., Waters, B., Schiebel, D., Young, W., & Golap, K. 2007, *ASP Conf. Ser.*, **376**, 127
- Nagao, T., Maiolino, R., De Breuck, C., et al. 2012, *A&A*, **542**, L34
- Papadopoulos, P. P., & Greve, T. R. 2004, *ApJ*, **615**, L29
- Planck Collaboration XIII. 2015, *A&A*, submitted [[arXiv:1502.01589](https://arxiv.org/abs/1502.01589)]
- Rawle, T. D., Egami, E., Bussmann, R. S., et al. 2014, *ApJ*, **783**, 59
- Riechers, D. A., Walter, F., Carilli, C. L., & Lewis, G. F. 2009, *ApJ*, **690**, 463
- Riechers, D. A., Bradford, C. M., Clements, D. L., et al. 2013, *Nature*, **496**, 329
- Rosenberg, M. J. F., van der Werf, P. P., Aalto, S., et al. 2015, *ApJ*, **801**, 72
- Scoville, N., Aussel, H., Sheth, K., et al. 2014, *ApJ*, **783**, 84
- Serjeant, S. 2012, *MNRAS*, **424**, 2429
- Smail, I., Ivison, R. J., & Blain, A. W. 1997, *ApJ*, **490**, L5
- Smolčić, V., Karim, A., Miettinen, O., et al. 2015, *A&A*, **576**, A127
- Sparre, M., Hayward, C. C., Springel, V., et al. 2015, *MNRAS*, **447**, 3548
- Spilker, J. S., Marrone, D. P., Aguirre, J. E., et al. 2014, *ApJ*, **785**, 149
- Spinoglio, L., Pereira-Santaella, M., Dasysa, K. M., et al. 2015, *ApJ*, **799**, 21
- Strandet, M., de Breuck, C., Vieira, J., Weiss, A., & SPT collaboration 2016, *ApJ*, submitted
- Vieira, J. D., Marrone, D. P., Chapman, S. C., et al. 2013, *Nature*, **495**, 344
- Weiß, A., Downes, D., Neri, R., et al. 2007, *A&A*, **467**, 955
- Weiß, A., De Breuck, C., Marrone, D. P., et al. 2013, *ApJ*, **767**, 88